

Minimal symmetry breaking of the cosmological principle and the Lemaître-Hubble diagram

André Tilquin, Galliano Valent

- Surface of the Earth \leftrightarrow universe
- Minimal symmetry breaking of the cosmological principle with comoving dust \Rightarrow axial Bianchi IX universes
- Lemaître-Hubble diagram of axial Bianchi IX universes
- Future experimental tests of axial Bianchi IX universes

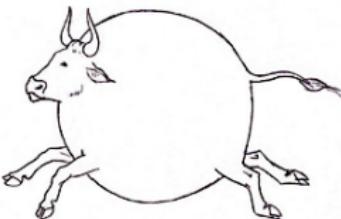
To the memory of Vaughan Jones

Symmetries reduce the complexity of systems and may lead to useful approximations. Observed deviations from the symmetric approximation may destroy all symmetries or only some.

Symmetries reduce the complexity of systems and may lead to useful approximations. Observed deviations from the symmetric approximation may destroy all symmetries or only some.

Ex 2d: On our planet, the paradigm of the **spherical cow** is useful:

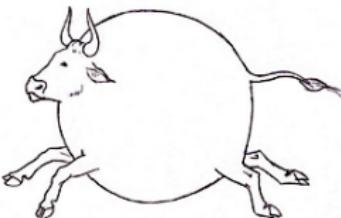
'geographic' symmetry breakings (easy to see): Mount Everest
8.8 km high, $8.8 \text{ km} \cdot 2\pi / (40\,000 \text{ km}) \approx 1.4 \cdot 10^{-3}$;



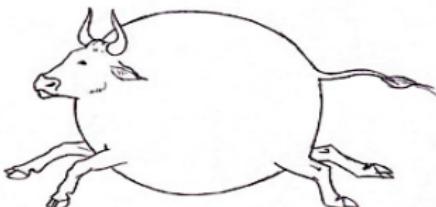
Symmetries reduce the complexity of systems and may lead to useful approximations. Observed deviations from the symmetric approximation may destroy all symmetries or only some.

Ex 2d: On our planet, the paradigm of the **spherical cow** is useful:

'geographic' symmetry breakings (easy to see): Mount Everest
8.8 km high, $8.8 \text{ km} \cdot 2\pi / (40000 \text{ km}) \approx 1.4 \cdot 10^{-3}$;



'geometric' sym. break. (less evident): (equatorial – polar) radius = 21.3 km, $21.3 \text{ km} \cdot 2\pi / (40000 \text{ km}) \approx 3.3 \cdot 10^{-3}$.



Example 3d: In our universe, the spherical cow is called *cosmological principle* and is useful:

geographic symmetry breakings: well established in the **Cosmic Microwave Background (CMB)**, at $\approx 10^{-5}$;

Example 3d: In our universe, the spherical cow is called *cosmological principle* and is useful:
geographic symmetry breakings: well established in the **Cosmic Microwave Background (CMB)**, at $\approx 10^{-5}$;
geometric symmetry breakings, modelled by axial Bianchi I universes,

$$d\tau^2 = dt^2 - a(t)^2 [dx^2 + dy^2] - c(t)^2 dz^2,$$

are signaled in CMB data by Cea [2014] at $\approx 10^{-10}$ (1σ),

Example 3d: In our universe, the spherical cow is called *cosmological principle* and is useful:

geographic symmetry breakings: well established in the **Cosmic Microwave Background (CMB)**, at $\approx 10^{-5}$;
geometric symmetry breakings, modelled by axial Bianchi I universes,

$$d\tau^2 = dt^2 - a(t)^2 [dx^2 + dy^2] - c(t)^2 dz^2,$$

are signaled in CMB data by Cea [2014] at $\approx 10^{-10}$ (1σ),
and in the Lemaître-Hubble diagram (740 type 1a supernovae,
redshift $z \leq 1.3$) by Tilquin, S. & Valent [2014] at $\approx 10^{-2}$ (1σ).

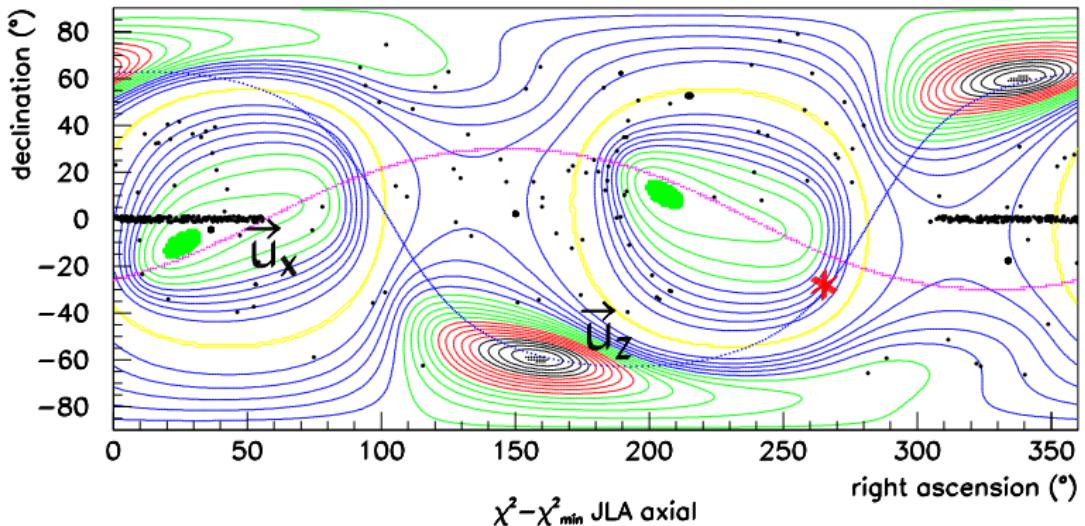


Figure: Black points represent 740 supernova positions. Note the accumulation of supernovae in the equatorial plane of the Earth and the absence of supernovae in the galactic plane (blue line). The red star is the direction towards our galactic center. Confidence level contours of privileged directions in arbitrary color codes for axial Bianchi I universes; best fit: \vec{U}_z .

Example 3d: In our universe, the spherical cow is called *cosmological principle* and is useful:
geographic symmetry breakings: well established in the **Cosmic Microwave Background (CMB)**, at $\approx 10^{-5}$;
geometric symmetry breakings, modelled by axial Bianchi I universes,

$$d\tau^2 = dt^2 - a(t)^2 [dx^2 + dy^2] - c(t)^2 dz^2,$$

are signaled in CMB data by Cea [2014] at $\approx 10^{-10}$ (1σ),
and in the Lemaître-Hubble diagram (740 type 1a supernovae,
redshift $z \leq 1.3$) by Tilquin, S. & Valent [2014] at $\approx 10^{-2}$ (1σ).

Again in 2014 Darling used tri-axial Bianchi I universes,

$$d\tau^2 = dt^2 - a(t)^2 dx^2 - b(t)^2 dy^2 - c(t)^2 dz^2,$$

Lemaître [1933], to fit the drift of 429 extra-galactic radio sources.

surface of the Earth	universe	sym. break.
sphericity*	cosmological principle	
static and isolated	Lemaître-Hubble flow	
mountains	over-densities	geographic
continental drift, variable duration of the day	peculiar velocities	geographic
constant rotation, oblate ellipticity	Bianchi I	geometric
precession period = 26 ky	position drift	geometric

*Flat-Earthers believe that sphericity is a conspiracy.

They must love flat universes too.

Future experimental tests

One year of Vera Rubin Observatory data should yield 50 000 supernovae (starting 2024 ?) and a precision better than

$$6 \cdot 10^{-4} ?$$

Future experimental tests

One year of Vera Rubin Observatory data should yield 50 000 supernovae (starting 2024 ?) and a precision better than

$$6 \cdot 10^{-4} ?$$

Complementarily, the James Webb Space Telescope is expected to observe some 200 type 1a supernovae up to redshift $z = 6$ in the next six years.

Future experimental tests

One year of Vera Rubin Observatory data should yield 50 000 supernovae (starting 2024 ?) and a precision better than

$$6 \cdot 10^{-4} ?$$

Complementarily, the James Webb Space Telescope is expected to observe some 200 type 1a supernovae up to redshift $z = 6$ in the next six years.

The Chinese Space Station Telescope should start operating in 2024 and observe some 1800 type 1a supernovae below a redshift of $z = 1.3$ in a time span of two years.

Future experimental tests

One year of Vera Rubin Observatory data should yield 50 000 supernovae (starting 2024 ?) and a precision better than

$$6 \cdot 10^{-4} ?$$

Complementarily, the James Webb Space Telescope is expected to observe some 200 type 1a supernovae up to redshift $z = 6$ in the next six years.

The Chinese Space Station Telescope should start operating in 2024 and observe some 1800 type 1a supernovae below a redshift of $z = 1.3$ in a time span of two years.

Exciting years ahead, inviting us to propose already now finer mathematical models to be tested.

A few mathematical facts on maximal symmetry and its minimal breakings

Theorem: The isometry group of a d -dimensional space or spacetime (with $d \geq 2$) is a Lie group of dimension $n \leq d(d+1)/2$.

Examples in $d = 2$:

The sphere has the $n = d(d+1)/2 = 3$ dimensional isometry group $O(3)$, it is “maximally symmetric”;

oblate and prolate axial ellipsoids (pumpkin and rugby ball) have the $n = 1$ dimensional isometry group $O(2)$;

generic ellipsoids have a discrete isometry group only, $n = 0$.

A few mathematical facts on maximal symmetry and its minimal breakings

Theorem: The isometry group of a d -dimensional space or spacetime (with $d \geq 2$) is a Lie group of dimension $n \leq d(d+1)/2$.

Examples in $d = 2$:

The sphere has the $n = d(d+1)/2 = 3$ dimensional isometry group $O(3)$, it is “maximally symmetric”;

oblate and prolate axial ellipsoids (pumpkin and rugby ball) have the $n = 1$ dimensional isometry group $O(2)$;

generic ellipsoids have a discrete isometry group only, $n = 0$.

Theorem (Guido Fubini 1903): The isometry group of a d -dimensional space, $d \geq 3$, cannot be of dimension $n = d(d+1)/2 - 1$.

A few mathematical facts on maximal symmetry and its minimal breakings

Theorem: The isometry group of a d -dimensional space or spacetime (with $d \geq 2$) is a Lie group of dimension $n \leq d(d+1)/2$.

Examples in $d = 2$:

The sphere has the $n = d(d+1)/2 = 3$ dimensional isometry group $O(3)$, it is “maximally symmetric”;

oblate and prolate axial ellipsoids (pumpkin and rugby ball) have the $n = 1$ dimensional isometry group $O(2)$;

generic ellipsoids have a discrete isometry group only, $n = 0$.

Theorem (Guido Fubini 1903): The isometry group of a d -dimensional space, $d \geq 3$, cannot be of dimension $n = d(d+1)/2 - 1$.

In $d = 2$ dimensions the cylinder is a counter example, because it has two isometries.

If we say: axial ellipsoids realize **minimal symmetry breakings** of the sphere, we mean:

- (1) These ellipsoids can be infinitesimally close to the sphere.
- (2) They have the highest possible number of symmetries:

$$n = d(d+1)/2 - 2 = 1.$$

If we say: axial ellipsoids realize **minimal symmetry breakings** of the sphere, we mean:

- (1) These ellipsoids can be infinitesimally close to the sphere.
- (2) They have the highest possible number of symmetries:

$$n = d(d+1)/2 - 2 = 1.$$

Examples in $d = 1 + 3$:

In relativistic cosmology, we are tempted to start with a maximally symmetric space-time: de Sitter spaces, Minkowski space or anti de Sitter spaces. However none of them admits dynamics and we must be more modest: $d = 3$.

If we say: axial ellipsoids realize **minimal symmetry breakings** of the sphere, we mean:

- (1) These ellipsoids can be infinitesimally close to the sphere.
- (2) They have the highest possible number of symmetries:

$$n = d(d+1)/2 - 2 = 1.$$

Examples in $d = 1 + 3$:

In relativistic cosmology, we are tempted to start with a maximally symmetric space-time: de Sitter spaces, Minkowski space or anti de Sitter spaces. However none of them admits dynamics and we must be more modest: $d = 3$.

The cosmological principle postulates maximally symmetric spaces of simultaneity: 3-spheres, \mathbb{R}^3 and pseudo 3-spheres with the $n = 6$ dimensional isometry groups: $O(4)$, $O(3) \ltimes \mathbb{R}^3$, $O(3, 1)$.

If we say: axial ellipsoids realize **minimal symmetry breakings** of the sphere, we mean:

- (1) These ellipsoids can be infinitesimally close to the sphere.
- (2) They have the highest possible number of symmetries:

$$n = d(d+1)/2 - 2 = 1.$$

Examples in $d = 1 + 3$:

In relativistic cosmology, we are tempted to start with a maximally symmetric space-time: de Sitter spaces, Minkowski space or anti de Sitter spaces. However none of them admits dynamics and we must be more modest: $d = 3$.

The cosmological principle postulates maximally symmetric spaces of simultaneity: 3-spheres, \mathbb{R}^3 and pseudo 3-spheres with the $n = 6$ dimensional isometry groups: $O(4)$, $O(3) \ltimes \mathbb{R}^3$, $O(3, 1)$.

Adding time as an orthogonal \mathbb{R} to these 3-spaces of simultaneity, one obtains the ‘Robertson-Walker’ universes.

Definition: A **minimal symmetry breaking** of the cosmological principle is

- (1) a smooth family of **deformations** of a maximally symmetric 3-space,
- (2) such that the isometry group of all deformations has maximal dimension, $n = 3(3 + 1)/2 - 2 = 4$ according to Fubini.

Definition: A minimal symmetry breaking of the cosmological principle is

- (1) a smooth family of **deformations** of a maximally symmetric 3-space,
- (2) such that the isometry group of all deformations has maximal dimension, $n = 3(3+1)/2 - 2 = 4$ according to Fubini.

Example: axial Bianchi I universes, 2 scale factors:

$$d\tau^2 = dt^2 - a^2 [dx^2 + dy^2] - c^2 dz^2,$$

3 translations + 1 rotation (around the z axis) = 4 symmetries.

Definition: A minimal symmetry breaking of the cosmological principle is

- (1) a smooth family of **deformations** of a maximally symmetric 3-space,
- (2) such that the isometry group of all deformations has maximal dimension, $n = 3(3+1)/2 - 2 = 4$ according to Fubini.

Example: axial Bianchi I universes, 2 scale factors:

$$d\tau^2 = dt^2 - a^2 [dx^2 + dy^2] - c^2 dz^2,$$

3 translations + 1 rotation (around the z axis) = 4 symmetries.

Counter-example: tri-axial Bianchi I universes, 3 scalefactors:

$$d\tau^2 = dt^2 - a^2 dx^2 - b^2 dy^2 - c^2 dz^2,$$

3 translations + no rotation = 3 symmetries.

Three symmetries: Bianchi universes

In 1898 Luigi Bianchi classifies all 3-dimensional, real Lie algebras. His list starts with seven 3-dimensional Lie algebras: Bianchi I, II, III, IV, V, VIII and IX.

Three symmetries: Bianchi universes

There is only one 1-dimensional Lie algebra: the Abelian one.

In 1898 Luigi Bianchi classifies all 3-dimensional, real Lie algebras.
His list starts with seven 3-dimensional Lie algebras: Bianchi I, II,
III, IV, V, VIII and IX.

Three symmetries: Bianchi universes

There is only one 1-dimensional Lie algebra: the Abelian one.

There are two 2-dimensional Lie algebras: the Abelian one and the (solvable) one of 2×2 triangular matrices with vanishing trace.

In 1898 Luigi Bianchi classifies all 3-dimensional, real Lie algebras. His list starts with seven 3-dimensional Lie algebras: Bianchi I, II, III, IV, V, VIII and IX.

Three symmetries: Bianchi universes

There is only one 1-dimensional Lie algebra: the Abelian one.

There are two 2-dimensional Lie algebras: the Abelian one and the (solvable) one of 2×2 triangular matrices with vanishing trace.

In 1898 Luigi Bianchi classifies all 3-dimensional, real Lie algebras. His list starts with seven 3-dimensional Lie algebras: Bianchi I, II, III, IV, V, VIII and IX. Bianchi I is Abelian (3 translations),

Three symmetries: Bianchi universes

There is only one 1-dimensional Lie algebra: the Abelian one.

There are two 2-dimensional Lie algebras: the Abelian one and the (solvable) one of 2×2 triangular matrices with vanishing trace.

In 1898 Luigi Bianchi classifies all 3-dimensional, real Lie algebras. His list starts with seven 3-dimensional Lie algebras: Bianchi I, II, III, IV, V, VIII and IX. Bianchi I is Abelian (3 translations), Bianchi II is the Heisenberg algebra. Bianchi III is the direct sum of the 1-dimensional and the solvable 2-dimensional Lie algebras. Bianchi IX is $so(3)$.

Three symmetries: Bianchi universes

There is only one 1-dimensional Lie algebra: the Abelian one.

There are two 2-dimensional Lie algebras: the Abelian one and the (solvable) one of 2×2 triangular matrices with vanishing trace.

In 1898 Luigi Bianchi classifies all 3-dimensional, real Lie algebras. His list starts with seven 3-dimensional Lie algebras: Bianchi I, II, III, IV, V, VIII and IX. Bianchi I is Abelian (3 translations), Bianchi II is the Heisenberg algebra. Bianchi III is the direct sum of the 1-dimensional and the solvable 2-dimensional Lie algebras. Bianchi IX is $so(3)$. In addition he finds two uncountable families of 3-dimensional Lie algebras, each indexed by a real parameter h : Bianchi VI $_h$, $h \neq 0, h \neq 1$ and VII $_h$, $h \geq 0$. ($h \neq$ Planck's constant!)

Three symmetries: Bianchi universes

There is only one 1-dimensional Lie algebra: the Abelian one.

There are two 2-dimensional Lie algebras: the Abelian one and the (solvable) one of 2×2 triangular matrices with vanishing trace.

In 1898 Luigi Bianchi classifies all 3-dimensional, real Lie algebras. His list starts with seven 3-dimensional Lie algebras: Bianchi I, II, III, IV, V, VIII and IX. Bianchi I is Abelian (3 translations), Bianchi II is the Heisenberg algebra. Bianchi III is the direct sum of the 1-dimensional and the solvable 2-dimensional Lie algebras. Bianchi IX is $so(3)$. In addition he finds two uncountable families of 3-dimensional Lie algebras, each indexed by a real parameter h : Bianchi VI $_h$, $h \neq 0, h \neq 1$ and VII $_h$, $h \geq 0$. ($h \neq$ Planck's constant!)

Bianchi also shows that all of these Lie algebras can be represented as **infinitesimal** isometries ('Killing vectors') on 3-spaces. Adding time as an orthogonal \mathbb{R} , one obtains the Bianchi universes.

Four symmetries: axial Bianchi universes

Bianchi classifies all 4-dimensional Lie algebras that occur as infinitesimal isometries on 3-spaces.

Four symmetries: axial Bianchi universes

Bianchi classifies all 4-dimensional Lie algebras that occur as infinitesimal isometries on 3-spaces.

He finds that all of these 4-dimensional Lie algebras contain 3-dimensional Lie sub-algebras. (Today we know that any real or complex 4-dimensional Lie algebra has a 3-dimensional ideal.) Therefore these 3-spaces give rise to universes, that are special types of Bianchi universes. Let us call them '**axial**' Bianchi universes.

Minimal symmetry breaking of the cosmological principle

Which of the Bianchi universes qualify as minimal symmetry breaking of the cosmological principle?

Minimal symmetry breaking of the cosmological principle

Which of the Bianchi universes qualify as minimal symmetry breaking of the cosmological principle?

Condition (1) ('Infinitesimally close to maximal symmetry')
eliminates the Bianchi II, III, IV, VI_h, and Bianchi VIII universes.

Minimal symmetry breaking of the cosmological principle

Which of the Bianchi universes qualify as minimal symmetry breaking of the cosmological principle?

Condition (1) ('Infinitesimally close to maximal symmetry')
eliminates the Bianchi II, III, IV, VI_h , and Bianchi VIII universes.

Condition (2) ('Four symmetries')
eliminates the axial Bianchi VII_h universes,
because they are isomorphic
to the axial Bianchi I universe for $h = 0$ and
to the axial Bianchi V universe for $h \neq 0$.

Minimal symmetry breaking of the cosmological principle

Which of the Bianchi universes qualify as minimal symmetry breaking of the cosmological principle?

Condition (1) ('Infinitesimally close to maximal symmetry')

eliminates the Bianchi II, III, IV, VI_h , and Bianchi VIII universes.

Condition (2) ('Four symmetries')

eliminates the axial Bianchi VII_h universes,

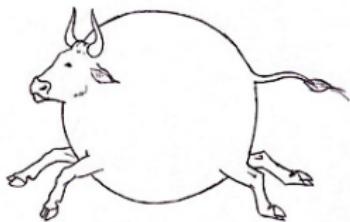
because they are isomorphic

to the axial Bianchi I universe for $h = 0$ and

to the axial Bianchi V universe for $h \neq 0$.

We remain with the axial Bianchi I, V and IX universes. They are smooth deformations of the Robertson-Walker universes with zero, negative and positive curvatures.

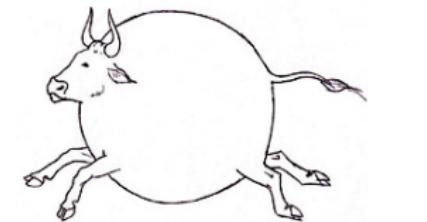
In pictures



Robertson-Walker

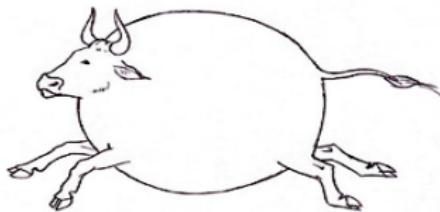
+ CMB + ...

In pictures



Robertson-Walker

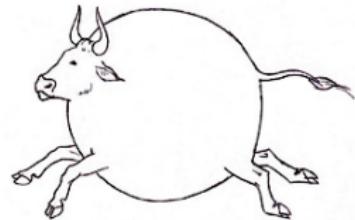
+ CMB + ...



axial Bianchi I, V, IX

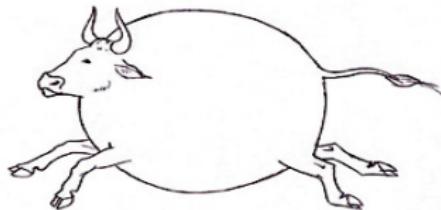
+ CMB + ...

In pictures



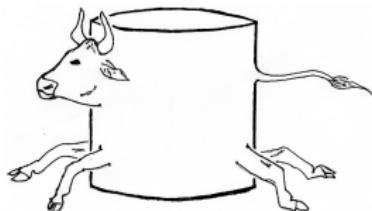
Robertson-Walker

+ CMB + ...



axial Bianchi I, V, IX

+ CMB + ...



Bianchi II, III, IV, VI_h, VIII + CMB + ...

Axial Bianchi V universes with comoving dust are incompatible with Einstein's equations.

1967 Farnsworth

Axial Bianchi V universes with comoving dust are incompatible with Einstein's equations.

In 1967 Farnsworth solved the Einstein equations for **axial** Bianchi V universes with **not necessarily comoving** dust and finds that none of his solutions supports comoving dust (except the Friedman solution).

Axial Bianchi V universes with comoving dust are incompatible with Einstein's equations.

In 1967 Farnsworth solved the Einstein equations for **axial** Bianchi V universes with **not necessarily comoving** dust and finds that none of his solutions supports comoving dust (except the Friedman solution).

In 2022 Galliano solved the Einstein equations for **not necessarily axial** Bianchi V universes with **comoving** dust and finds that none of his solutions supports the axial symmetry (except the Friedman solution).

Axial Bianchi V universes with comoving dust are incompatible with Einstein's equations.

In 1967 Farnsworth solved the Einstein equations for **axial** Bianchi V universes with **not necessarily comoving** dust and finds that none of his solutions supports comoving dust (except the Friedman solution).

In 2022 Galliano solved the Einstein equations for **not necessarily axial** Bianchi V universes with **comoving** dust and finds that none of his solutions supports the axial symmetry (except the Friedman solution).

In 2021 Akarsu, Di Valentino, Kumar, Ozyigit & Sharma considered **tri-axial** Bianchi V universes with $c^2 = ab$ and **comoving** dust. Galliano's bonus: exact solution of Einstein's equations. We also prove that the isometry group is 3-dimensional (except for $a = b$, the Friedman solution).

Axial Bianchi IX universes with comoving dust

Axial Bianchi IX universes with comoving dust

- Two scale factors: $a(t) = b(t)$ and $c(t)$

Axial Bianchi IX universes with comoving dust

- Two scale factors: $a(t) = b(t)$ and $c(t)$

that we linearize:

$$a = a_F [1 - \frac{1}{2} \eta], \quad c = a_F [1 + \eta], \quad |\eta(t)| \ll 1,$$

(\cdot_F for Friedman).

Axial Bianchi IX universes with comoving dust

- Two scale factors: $a(t) = b(t)$ and $c(t)$

that we linearize:

$$a = a_F [1 - \frac{1}{2} \eta], \quad c = a_F [1 + \eta], \quad |\eta(t)| \ll 1,$$

(\cdot_F for Friedman).

- Seven parameters (\cdot_0 for evaluation *today* = t_0):
three for the underlying spherical Friedman universe

$$H_{F0}, \Omega_{\Lambda 0}, \Omega_{m0},$$

two for the direction of the axial symmetry axis

(right ascension, declination),

and two for ‘ellipticity’ η_0 and ‘Hubble stretch’ η'_0 , ($' := d/dt$).

Axial Bianchi IX universes with comoving dust

- Two scale factors: $a(t) = b(t)$ and $c(t)$

that we linearize:

$$a = a_F [1 - \frac{1}{2} \eta], \quad c = a_F [1 + \eta], \quad |\eta(t)| \ll 1,$$

(\cdot_F for Friedman).

- Seven parameters (\cdot_0 for evaluation *today* = t_0):
three for the underlying spherical Friedman universe

$$H_{F0}, \Omega_{\Lambda 0}, \Omega_{m0},$$

two for the direction of the axial symmetry axis

(right ascension, declination),

and two for ‘ellipticity’ η_0 and ‘Hubble stretch’ η'_0 , ($' := d/dt$).

- Axial Bianchi I is the five-parameter sub-model with vanishing curvature $\Omega_{\Lambda 0} + \Omega_{m0} = 1$ and vanishing ellipticity $\eta_0 = 0$.

Consider a supernova of absolute luminosity L . It emits light at cosmic time t_e . The light arrives today, t_0 , and here with a redshift z , an apparent luminosity ℓ and at an angle θ with respect to the axial symmetry axis.

Consider a supernova of absolute luminosity L . It emits light at cosmic time t_e . The light arrives today, t_0 , and here with a redshift z , an apparent luminosity ℓ and at an angle θ with respect to the axial symmetry axis. From the [kinematics](#) we obtain:

$$z + 1 \sim \frac{a_{F0}}{a_{Fe}} \left[1 - \frac{1 - 3 \cos^2 \theta}{2} (\eta_0 - \eta_e) \right], \quad (0)$$

$$\ell \sim \ell_F \left[1 + \frac{1 - 3 \cos^2 \theta}{2} \left(\eta_0 - 5\eta_e + 4 \frac{\chi}{\tan \chi} \bar{\eta} \right) \right], \quad (1)$$

with \cdot_e standing for evaluation at t_e and with

$$\ell_F = \frac{L}{4\pi a_{F0}^2 \sin^2 \chi} \left(\frac{a_{Fe}}{a_{F0}} \right)^2, \quad \chi := \int_{t_e}^{t_0} \frac{1}{a_F}, \quad \bar{\eta} := \frac{1}{\chi} \int_{t_e}^{t_0} \frac{\eta}{a_F}.$$

Consider a supernova of absolute luminosity L . It emits light at cosmic time t_e . The light arrives today, t_0 , and here with a redshift z , an apparent luminosity ℓ and at an angle θ with respect to the axial symmetry axis. From the [kinematics](#) we obtain:

$$z + 1 \sim \frac{a_{F0}}{a_{Fe}} \left[1 - \frac{1 - 3 \cos^2 \theta}{2} (\eta_0 - \eta_e) \right], \quad (0)$$

$$\ell \sim \ell_F \left[1 + \frac{1 - 3 \cos^2 \theta}{2} \left(\eta_0 - 5\eta_e + 4 \frac{\chi}{\tan \chi} \bar{\eta} \right) \right], \quad (1)$$

with \cdot_e standing for evaluation at t_e and with

$$\ell_F = \frac{L}{4\pi a_{F0}^2 \sin^2 \chi} \left(\frac{a_{Fe}}{a_{F0}} \right)^2, \quad \chi := \int_{t_e}^{t_0} \frac{1}{a_F}, \quad \bar{\eta} := \frac{1}{\chi} \int_{t_e}^{t_0} \frac{\eta}{a_F}.$$

In the limit of zero curvature, $\sin \chi \rightarrow \chi$ and $\chi / \tan \chi \rightarrow 1$ hold, and we recover our 2014 results for axial Bianchi I.

The **dynamics** is given by Einstein's equations:

- $a_F(t)$ satisfies Friedman's equations with final conditions:
 $a_F(t_0) = a_{F0}$ and $a'_F(t_0) = a_{F0} H_{F0}$.
- $\eta(t)$ satisfies

$$\eta'' + 3 H_F \eta' + 8 \frac{\eta}{a_F^2} \sim 0, \quad (2)$$

which admits a unique solution with final conditions:

ellipticity η_0 and Hubble stretch η'_0 .

The **dynamics** is given by Einstein's equations:

- $a_F(t)$ satisfies Friedman's equations with final conditions:
 $a_F(t_0) = a_{F0}$ and $a'_F(t_0) = a_{F0} H_{F0}$.
- $\eta(t)$ satisfies

$$\eta'' + 3 H_F \eta' + 8 \frac{\eta}{a_F^2} \sim 0, \quad (2)$$

which admits a unique solution with final conditions:

ellipticity η_0 and Hubble stretch η'_0 .

In the limit of zero curvature, the last term of equation (2) vanishes, and we recover our 2014 result for axial Bianchi I.

The **dynamics** is given by Einstein's equations:

- $a_F(t)$ satisfies Friedman's equations with final conditions:
 $a_F(t_0) = a_{F0}$ and $a'_F(t_0) = a_{F0} H_{F0}$.
- $\eta(t)$ satisfies

$$\eta'' + 3 H_F \eta' + 8 \frac{\eta}{a_F^2} \sim 0, \quad (2)$$

which admits a unique solution with final conditions:

ellipticity η_0 and Hubble stretch η'_0 .

In the limit of zero curvature, the last term of equation (2) vanishes, and we recover our 2014 result for axial Bianchi I.

Finally the emission time t_e is eliminated in favour of the redshift by inverting $z(t_e)$, equation (0). Then the apparent luminosity is computed and compared to the observed one and the seven initial parameters are varied in order to optimize the fit for all observed supernovae.

Conclusions

In 2014 we confronted the 5-parameter axial Bianchi I with the 740 supernovae up to redshift 1.3 and found only a $1-\sigma$ signal.

Therefore we think that a fit of the 7-parameter axial Bianchi IX model to the Lemaître-Hubble diagram of type 1a supernovae becomes reasonable only once we can include the data expected from the Vera Rubin Observatory, the James Webb Space Telescope, the Chinese Space Station Telescope, ...

Conclusions

In 2014 we confronted the 5-parameter axial Bianchi I with the 740 supernovae up to redshift 1.3 and found only a $1-\sigma$ signal.

Therefore we think that a fit of the 7-parameter axial Bianchi IX model to the Lemaître-Hubble diagram of type 1a supernovae becomes reasonable only once we can include the data expected from the Vera Rubin Observatory, the James Webb Space Telescope, the Chinese Space Station Telescope, ...

In the same year 2014 Cea fitted the axial Bianchi I model to the WMAP and Planck data at redshift 1090. He also finds a $1-\sigma$ signal. Although his Hubble stretch has opposite sign and is smaller than ours by eight orders of magnitude, our results are compatible with his.

Again in 2014 Darling used the tri-axial Bianchi I model (7 parameters) to fit the drift of 429 extra-galactic radio sources measured by Titov & Lambert in 2013 using Very Long Baseline Interferometry. His main Hubble stretch has the same sign as ours but is ten time larger and the results are again compatible statistically.

Again in 2014 Darling used the tri-axial Bianchi I model (7 parameters) to fit the drift of 429 extra-galactic radio sources measured by Titov & Lambert in 2013 using Very Long Baseline Interferometry. His main Hubble stretch has the same sign as ours but is ten time larger and the results are again compatible statistically.

- Waiting and preparing for the promised promising data of type 1a supernovae, a [combined analysis](#) of axial Bianchi IX universes with type 1a supernovae, Cosmic Microwave Background, drift of radio sources and Baryonic Acoustic Oscillations (and maybe weak lensing or black-hole mergers) is called for [now](#).

Traudi of Oberwolfach

To illustrate the cosmological principle here is a story set in the Black Forest. Next to the Oberwolfach Research Institute for Mathematics there is a farm, home to Traudi, her farmer's preferred cow. Traudi is ill and the veterinarian helpless. Sparing no effort, the farmer calls on a physician, more expensive, but as helpless as his colleague. The farmer has a nephew with a PhD in biology and asks him to see Traudi, again without success. Finally, when he sees a theoretical physicist on his way to a conference, the farmer, driven to despair, asks him for help. The physicist sits down next to Traudi, pulls out his note pad and starts calculating. During hours the farmer watches the physicist's intense concentration from a respectful distance and feels a timid ripple of optimism. He pulls closer, caresses Traudi between the horns and asks: 'Is there any hope?' 'Indeed there is', replies the physicist with unconcealed pride, 'I just solved the case of the spherical cow.'